The Anisotropy Search Program For the Pierre Auger Observatory

Roger Clay,[1] for the Pierre Auger Collaboration
(1) Department of Physics, University of Adelaide, S. Australia 5005

Abstract

The Pierre Auger Southern Observatory in Argentina has begun taking data as it is being developed up to a final enclosed area of 3000 square kilometres. A key aspect of the project is to provide information on the origin of the highest energy cosmic rays through understanding the arrival directions of those particles. To avoid claims of a spurious anisotropy detection because trials have not been properly accounted, the Auger Collaboration has agreed '*a priori*' to the analysis prescription presented here. It specifies the following:

1. The accumulation time for the (future) data set to be analyzed.
2. The anisotropy 'targets', each assigned a chance probability level.
3. The analysis procedure for each trial.

A positive result will be claimed for any target search only if its chance probability is less than its assigned level. The levels are chosen so that the total chance probability for one or more positive results is 0.001.

Exploratory searches beyond this prescription will be encouraged, but the Auger Collaboration will not assign any confidence level to anisotropies that may be discovered that way. Any such discovery would identify a good target for a prescription to be used with a subsequent Auger data set.

1. Introduction: The Need for a Prescription

A major aim of the Pierre Auger Collaboration is the identification of deviations from isotropy (the anisotropy) in the flux of cosmic rays. This aim is based on the expectation that rigidity dependent propagation will begin to allow a form of directional astronomy with cosmic rays at the highest energies. In past studies, the inevitable statistical noise in the data, and subtle time variations in the array stability, have caused some claims to be made for observed anisotropies which have not been confirmed. This has partially resulted from an inadequate knowledge of the number of trials involved in the analysis. The Auger project aims to avoid the latter uncertainties by an *a priori* definition of the analysis process so that the chance probability of any positive detection can be rigorously evaluated. We will refer to this as an *a priori* analysis prescription.

A prescription for an anisotropy search specifies the search procedure *a*

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priori so that a positive effect can be published with confidence in its statistical significance. It protects the limited number of important searches by documenting ahead of time the search that will be performed on a specified dataset.

This does not preclude a thorough exploration of the data for expected or unexpected anisotropies but such an exploration would not have immediate discovery potential. It would, however, make a case for the result of that exploration to be included in considering a prescription for a future dataset. We thus envisage that new Auger datasets will regularly become available as the array continues to operate and a specific and unique prescription will be determined *a priori* for each of those datasets.

2. The Prescription

2.1. Data Set Characterization

The Auger array is currently under development and is progressively increasing in collecting area. The early anisotropy data will be more complicated to analyse than those from a later period in which the array is more stable. However, we wish to test the prescription technique and its associated discipline as early as possible. We propose defining the start of our first dataset for analysis on 8 August 2003 (when the array will be well into its pre-production phase). The dataset will end on 16 May 2004. This ensures that the second prescription will relate to a much more stable period for the array. The event directions and energies used will be those produced by the official Auger reconstruction program in use at the time of the end of the dataset. We anticipate that this first dataset will contain about 10^4 events with energies above 1 EeV and about 300 events above 10 EeV. In the analysis of this first anisotropy data, we will use events with energies greater than 1 EeV.

2.2. The Targets

We propose that a publishable positive anisotropy result requires a chance probability (after accounting for statistical trials) that is less than 0.001. We thus propose to prioritize search targets by partitioning that overall 0.001 probability.

With our defined dataset, this prescription covers searches either over an area specified by a radial distance around a potential source direction, or in a specific (point source) direction. Our prescription is:

Our anisotropy search has a positive publishable result if one or more of the following sources have event numbers greater than required for the corresponding chance probabilities in the specified energy and angular ranges.

Source	*Search radius*	*Energy Range (log(eV))*	*Chance Prob.*
Galactic Centre	15°	All	.00035
Galactic Centre	Point Source	≤ 18.5	.00025
AGASA/SUGAR	Point Source	≤ 18.5	.00025
NGC0253	5°	≥ 19.5	.00005
NGC3256	5°	≥ 19.5	.00005
Centaurus A	5°	≥ 19.5	.00005

The AGASA experiment (Hayashida *et al.* 1999), with support from SUGAR data (Bellido *et al.* 2001), has identified a possible excess from either the galactic centre direction or a close by direction within a limited energy range. That direction has always been an *a priori* key source for the Auger Observatory and the first prescription source reflects this. The angular range specified covers the 'SUGAR direction' and also is intended to cover the expected source extension following propagation in the galactic magnetic field at characteristic Auger energies (Clay 2001). The second and third sources relate to those observations. NGC253 is a nearby (distance estimate 2.5Mpc) starburst galaxy which has been shown to emit a substantial flux of VHE gamma-ray photons (Itoh *et al.* 2002), NGC3256 is a nearby (33Mpc, Moran *et al.* 1999) merging multiple galaxy system experiencing highly luminous starburst, and Cen A is our nearest active galaxy at a distance estimated as 3.4Mpc (Bird and Clay 1990). The latter sources are outside the galactic plane and the angular range is intended to approximate to a point spread function for the (largely unknown) halo and intergalactic magnetic fields.

2.3. Evaluation Method

The final part of the analysis prescription is to specify how each of the above probabilities is to be evaluated. This initial data set will be small compared to data sets that will accumulate with the full Auger Observatory, so we cannot expect to have sensitivity to the small anisotropies that will eventually be measurable. Also, this first data set will be taken while the array is growing and changing shape. The analysis techniques prescribed here may be inappropriate for the better data sets of the future.

An ensemble of simulation data sets will be used to evaluate the probability that the number of arrival directions in a target solid angle would be as great or greater than the observed number. The simulation data sets are constructed from the actual data set by a 2-dimensional scrambling method. The method is therefore testing the null hypothesis that an excess from a certain region occurred as a statistical fluctuation of isotropy.

The scrambling method is not perfect. A large excess from one part of the sky will distort the sidereal time distribution that would pertain for isotropic

cosmic rays. The method is conservative, however, in the sense that the chance probabilities may be slightly overestimated, making the results appear less significant than they might actually be.

The method prescribed here is meant to be explicit. There is not space to justify each part of it. More complicated procedures might yield more reliable probability evaluations, but simplicity has been a goal. Here are the explicit rules:
* Use zenith angles in the range 0-70 degrees.
* Computing the on-source number of arrival directions: (1) For the simple solid angle targets, an arrival direction is in the target if it differs from the target's central direction by less than the specified radius. (2) For the point source use $3^\circ \times (3 * 10^{19}/E)$ as the r.m.s. of the particle deviation angle.
* Make 5 bins of zenith angles uniformly in cos(theta) between cos(theta)=1.0 and cos(theta)=0.5.
* To produce a simulation data set of N directions, sample N arrival directions from the real data set. For each one, keep the zenith angle, but sample a new sidereal time from events in the same zenith angle bin, and sample a new azimuth also from events in that zenith angle bin.
* If the assigned significance probability is P, do n=1000/P simulations in determining the chance probability of the observed excess. This gives an expected number of simulation data sets showing a greater positive effect than the real data (at the assigned probability) as nP=1000. The statistical uncertainty in the obtained probability would then be roughly 3%, i.e. sqrt(1000)/1000.
* These rules can be coded into evaluation programs of different languages and in different ways. Alternative analysis programs will be checked against each other using artificial data sets.

3. Summary

We have defined here a specific prescription for the directional analysis of the first pre-defined Auger data set. When that analysis is complete, a new prescription will be defined for the next data set and so on as the experiment continues to collect new data.

4. References

J.A. Bellido *et al.* Astroparticle Physics 15, 167 (2001)
Bird, D.J. and Clay, R.W., Proc. Astron. Soc. Aust. 8, 266, (1990)
G.L. Cassiday *et al.* Nucl. Phys. B (Proc. Suppl.) 14A, 291 (1990)
R.W. Clay, Pub. Astron. Soc. Aust. 18, No 2, 148, (2001)
N. Hayashida *et al.* Astroparticle Physics 10, 303 (1999)
C. Itoh *et al.* Astronomy and Astrophysics 396, L1 (2002)
E.C. Moran *et al.* Ap.J. 526, 649, (1999)